

INNER-SHELL $1s$ – $2p$ SOFT X-RAY ABSORPTION LINES

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ABSTRACT

The HULLAC atomic code is used to compute wavelengths and oscillator strengths for the $1s$ – np inner-shell absorption lines in Li-like to F-like ions of neon, magnesium, aluminum, silicon, sulfur, argon, calcium, and iron. Many of these lines are expected to be observed in *Chandra* and *XMM-Newton* high-resolution X-ray spectra of active galaxies. The new atomic data are incorporated in the ION code for spectral modeling of photoionized plasmas. The calculated spectra are subsequently compared with the spectrum of NGC 3783 and show very good agreement. The usefulness of these lines as diagnostics for the ionization state, column densities, and velocities in line-of-sight photoionized gas is highlighted.

Subject headings: atomic data — atomic processes — galaxies: active — galaxies: nuclei — line: formation — X-rays: galaxies

1. INTRODUCTION

The launch of the *Chandra* and *XMM-Newton* observatories along with the grating spectrometers on board has generated great interest in high-resolution X-ray spectroscopy of a wide variety of astrophysical X-ray sources, in particular active galactic nuclei (AGNs). The high-resolution grating observations of a handful of AGNs, e.g., NGC 5548 (Kaastra et al. 2000), NGC 3783 (Kaspi et al. 2000, 2001), MCG –6-30-15 (Branduardi-Raymont et al. 2001; Lee et al. 2001), and IRAS 13349+2438 (Sako et al. 2001) clearly show the presence of dozens of strong X-ray absorption lines originating from highly charged gas along the line of sight. While many of the absorption lines are due to transitions in H-like and He-like ions, some are clearly due to inner-shell transitions in lower ionization species. Perhaps the most prominent feature of this kind is the unresolved transition array (UTA) of inner-shell $n = 2$ to 3 (mainly $2p$ – $3d$) lines pertaining to various M-shell Fe ions identified in IRAS 13349+2438 (Sako et al. 2001). The atomic data needed to model this UTA feature have recently been published in Behar, Sako, & Kahn (2001). Another example is the inner-shell $1s$ – $2p$ feature of Li-like Si¹¹⁺ in NGC 3783 reported by Kaspi et al. (2001). The analogous $1s$ – $2p$ lines for Li-like O⁵⁺ have been calculated by Pradhan (2000) using the *R*-matrix method and identified in MCG –6-30-15 by Lee et al. (2001). Sako et al. (2002) have found evidence for additional $1s$ – np lines arising from lower charge states of oxygen (O⁴⁺ and O³⁺) in MCG –6-30-15. Recently, Nahar, Pradhan, & Zhang (2001) published results for Li-like C³⁺, O⁵⁺, and Fe²³⁺ obtained using the same method as Pradhan (2000).

Models such as those presented in Kaspi et al. (2001) predict that inner-shell transitions from the ground level of many more L-shell ions (Li-like to F-like) are indeed expected for X-ray illuminated photoionized plasmas. However, the available atomic data are few. The present paper aims at providing comprehensive results for the resonant

$1s$ – np transitions in Li-like through F-like ions of neon, magnesium, aluminum, silicon, sulfur, argon, calcium, and iron. For that purpose, we employ HULLAC (Hebrew University Lawrence Livermore Atomic Code), which uses the parametric potential approximation. This method is by far more efficient than the previously used methods. For the present case of inner-shell transitions in highly charged ions, it is also expected to be very accurate. In the following, § 2 describes the method of calculation and the new atomic data, and § 3 shows theoretical models for AGN spectra and a comparison with the spectrum of NGC 3783 obtained with the High Energy Transmission Grating Spectrometer (HETGS) on board *Chandra*.

2. ATOMIC CALCULATIONS

In order to calculate the wavelengths and oscillator strengths of interest, we use the multiconfiguration, relativistic HULLAC computer package (Bar-Shalom, Klapisch, & Oreg 2001). The level energies in HULLAC are calculated using the relativistic version of the parametric potential (PP) method by Klapisch et al. (1977). This relativistic approach is perfectly adequate for the high-energy inner-shell excitations calculated in the present work. The high efficiency of HULLAC allows us to calculate the numerous transitions involved relatively quickly. Unlike some other atomic methods, the PP approach is hardly hampered by the considerable complexity of the electronic structure in the ions considered. Calculations are carried out for $1s$ – np ($n \leq 7$) transitions in the L-shell ions (Li-like to F-like) of Ne, Mg, Al, Si, S, Ar, Ca, and Fe. The atomic structure calculations include the ground configurations $1s^2 2l^x$ ($x = 1$ to 7). For each ion, all of the $1s$ – np photoexcitation transitions from the ground level are calculated, among which the $1s$ – $2p$ are the strongest by far. Electric and magnetic dipole and quadrupole transitions are computed. The contributions, however, of excitations other than electric-dipole are negligible.

The calculated $1s$ – np wavelengths (λ_{ij}) and oscillator strengths (f_{ij}) for the strongest absorption lines in the He I to F I isoelectronic sequences are listed in Table 1. Only lines with $f_{ij} > 0.1$, or the strongest line for a given ion, are included in the table. These lines represent the major part

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TABLE 1
 STRONGEST $1s\text{-}np$ ABSORPTION LINES FROM THE GROUND LEVEL OF HE-LIKE TO F-LIKE IONS

ELEMENT	ISOELECTRONIC SEQUENCE	λ_{ij}		f_{ij}		A_{ji} (10^{13} s^{-1})	$\Sigma(A + A^a)$ (10^{13} s^{-1})	
		Present (\AA)	Previous (\AA) ^a	Present	Previous ^a			
Ne.....	He I	13.448	13.447	0.657	0.724	0.81	0.81	
	...	11.553 ^b	11.547	0.159	0.149	0.27	0.28	
	Li I	13.646	...	0.381	...	0.68	1.30	
	...	13.648	...	0.187	...	0.67	1.43	
	Be I	13.814	...	0.595	...	0.69	8.25	
	...	12.175 ^b	...	0.126	...	0.19	0.21	
	B I	14.020	...	0.190	...	0.64	3.55	
	...	14.047	...	0.166	...	0.28	18.4	
	C I	14.239	...	0.147	...	0.16	24.9	
	...	14.202	...	0.124	...	0.14	5.76	
	N I	14.371	...	0.119	...	0.26	27.3	
	O I	14.526	...	0.106	...	0.34	25.5	
	F I	14.631	...	0.062	...	0.39	0.58	
	Mg.....	He I	9.170	9.169	0.691	0.742, 0.745 ^c	1.81	1.81
		...	7.854 ^b	7.851	0.159	0.151, 0.152 ^c	0.57	0.61
Li I		9.281	...	0.402	...	1.56	2.12	
...		9.283	...	0.194	...	1.50	2.36	
Be I		9.378	...	0.631	...	1.60	9.67	
...		8.201 ^b	...	0.132	...	0.44	0.46	
B I		9.498	...	0.207	...	1.53	5.08	
...		9.514	...	0.188	...	0.69	20.8	
C I		9.631	...	0.166	...	0.40	29.1	
...		9.607	...	0.143	...	0.34	7.95	
N I		9.718	...	0.132	...	0.62	32.5	
O I		9.816	...	0.125	...	0.86	35.7	
F I		9.895	...	0.076	...	1.04	41.1	
Al.....		He I	7.757	7.757	0.693	0.750	2.56	2.56
		...	6.637 ^b	6.635	0.159	0.152, 0.153 ^c	0.80	0.85
	Li I	7.844	...	0.410	...	2.22	2.76	
	...	7.486	...	0.196	...	2.12	3.05	
	Be I	7.921	...	0.645	...	2.28	10.6	
	...	6.907 ^b	...	0.134	...	0.63	0.64	
	B I	8.016	...	0.214	...	2.22	6.19	
	...	8.029	...	0.198	...	1.03	21.9	
	C I	8.121	...	0.176	...	0.59	30.8	
	...	8.101	...	0.152	...	0.52	9.58	
	N I	8.193	...	0.137	...	0.91	34.8	
	O I	8.272	...	0.131	...	1.28	40.2	
	F I	8.337	...	0.082	...	1.57	47.7	
	Si.....	He I	6.648	6.648	0.701	0.757	3.53	3.53
		...	5.682 ^b	5.681	0.159	0.152	1.09	1.16
Li I		6.717	...	0.418	...	3.09	3.60	
...		6.719	...	0.197	...	2.90	3.92	
Be I		6.778	...	0.656	...	3.17	11.6	
...		5.896 ^b	...	0.136	...	0.87	0.89	
B I		6.854	...	0.221	...	3.14	7.58	
...		6.864	...	0.209	...	1.48	22.9	
C I		6.939	...	0.185	...	0.86	32.4	
...		6.923	...	0.161	...	0.75	11.6	
N I		6.999	...	0.141	...	1.28	37.1	
O I		7.063	...	0.137	...	1.83	44.6	
F I		7.119	...	0.086	...	2.26	53.7	
S.....		He I	5.039	5.039	0.711	0.767	6.23	6.23
		...	4.300 ^b	4.299	0.158	0.153	1.90	2.01
	Li I	5.084	...	0.432	...	5.57	6.01	
	...	5.086	...	0.196	...	5.05	6.32	
	Be I	5.126	...	0.672	...	5.68	14.5	
	...	4.441 ^b	...	0.139	...	1.56	1.58	
	B I	5.176	...	0.234	...	5.82	11.4	
	...	5.183	...	0.230	...	2.86	25.1	
	C I	5.234	...	0.208	...	1.69	35.7	
	...	5.222	...	0.175	...	1.42	17.0	

TABLE 1—Continued

ELEMENT	ISOELECTRONIC SEQUENCE	λ_{ij}		f_{ij}		A_{ji} (10^{13} s^{-1})	$\Sigma(A + A^a)$ (10^{13} s^{-1})
		Present (Å)	Previous (Å) ^a	Present	Previous ^a		
Ar	N I	5.276	...	0.148	...	2.36	41.2
	O I	5.320	...	0.145	...	3.42	52.8
	F I	5.359	...	0.093	...	4.30	65.5
	He I	3.949	3.949	0.717	0.775	10.2	10.2
	...	3.366 ^b	3.365	0.156	0.155	3.06	3.25
	Li I	3.981	...	0.444	...	9.34	9.69
	...	3.983	...	0.192	...	8.06	9.66
	Be I	4.010	...	0.681	...	9.42	18.5
	...	3.464 ^b	...	0.140	...	2.59	2.62
	B I	4.050	...	0.250	...	5.08	27.7
	...	4.045	...	0.244	...	9.95	16.9
	C I	4.086	...	0.236	...	3.15	39.2
	...	4.078	...	0.175	...	2.34	24.6
	N I	4.117	...	0.153	...	4.01	45.4
...	4.115	...	0.103	...	4.06	45.4	
Ca	O I	4.148	...	0.151	...	5.84	61.1
	F I	4.176	...	0.098	...	7.46	77.4
	He I	3.177	3.177	0.717	0.782	15.8	15.8
	...	2.706 ^b	2.705	0.154	0.155	4.68	4.97
	Li I	3.200	...	0.454	...	14.8	15.0
	...	3.202	...	0.184	...	12.0	14.0
	Be I	3.221	...	0.685	...	14.7	24.1
	...	2.776 ^b	...	0.140	...	4.03	4.08
	B I	3.250	...	0.267	...	8.43	31.1
	...	3.247	...	0.252	...	16.0	24.7
	C I	3.276	...	0.273	...	5.65	43.0
	...	3.270	...	0.158	...	3.29	34.6
	N I	3.301	...	0.156	...	6.38	50.1
	...	3.299	...	0.106	...	6.51	50.1
Fe	O I	3.323	...	0.154	...	9.28	70.0
	F I	3.344	...	0.102	...	12.1	90.5
	He I	1.850	1.851	0.689	0.798, 0.704 ^c	44.7	44.7
	...	1.573 ^b	1.575, 1.573 ^c	0.144	0.156, 0.138 ^c	13.0	13.8
	Li I	1.861	1.860 ^d	0.469	0.491 ^d	45.1	45.2
	...	1.864	1.864 ^d	0.147	0.146 ^d	28.3	32.4
	Be I	1.870	...	0.666	...	42.3	52.6
	...	1.605 ^b	...	0.134	...	11.6	11.8
	B I	1.883	...	0.300	...	28.2	52.6
	...	1.883	...	0.259	...	48.7	65.2
	C I	1.895	...	0.397	...	24.6	64.4
	N I	1.908	...	0.154	...	18.8	69.7
	...	1.906	...	0.119	...	21.8	72.1
	O I	1.918	...	0.151	...	27.4	106.2
F I	1.927	...	0.109	...	21.2	144.0	

^a Primarily Verner, Verner, & Ferland 1996.

^b $1s-3p$ transitions.

^c NIST database (http://physics.nist.gov/cgi-bin/AtData/main_asd); quoted only where available and different from Verner, Verner, & Ferland (1996).

^d Nahar, Pradhan, & Zhang 2001.

(>70%, and in most cases >90%) of the absorption effect for the respective transitions. Thus, they should be sufficient for the analysis of most astrophysical spectra. We note, however, that the ION code, for which results are shown below, does include more lines than those presented in Table 1. The strong lines are predominantly $1s-2p$ lines, but also $1s-3p$ lines in the He I and Be I sequences (see Table 1). More specifically, the strongest lines usually arise for $1s-2p_{3/2}$ transitions. The Einstein coefficient (A_{ji}), as well as the total rate for depletion of the inner-shell excited level [$\Sigma(A + A^a)$], which includes both radiative decay (A) and autoionization

(A^a) rates, are also given in Table 1. The total depletion rate is needed to compute the fluorescence yield $A_{ji}/\Sigma(A + A^a)$ and natural width associated with each absorption line.

As can be seen in Table 1, and as expected, the oscillator strength depends only very weakly on the element and is much more contingent upon the isoelectronic sequence. Table 1 shows that for each element, the oscillator strength generally increases with the ionic charge, as more vacancies become available in the $2p$ subshell. This trend appears to be violated by the lines of the Be-like ions, which feature a $1s-2p$ line that is stronger than the lines of the Li-like ions.

However, the total oscillator strength, summed over all of the $1s-2p$ excitations for Li-like and Be-like species, is in fact comparable, as expected for these systems, both of which have the maximum possible six vacancies in the $2p$ subshell. The apparent dominance of the Be I sequence is only due to the fact that for the Be-like ions, the entire oscillator strength is concentrated in one $1s-2p_{3/2}$ line. The same stands for the $1s-3p_{3/2}$ transition (as well as for higher n), which explains why those transitions in Be-like ions are strong enough ($f_{ij} > 0.1$) to make it into the table. Conversely, for the Li-like ions, the total $1s-2p_{3/2}$ line strength is distributed among several lines.

The accuracy of all of the data presented here is expected to be very high, but owing to the relativistic method in HULLAC, the data become increasingly accurate for higher charge states. In order to give a rough idea of the uncertainties that could be expected from these atomic values, we include in Table 1, in addition to the new inner-shell excited lines, the leading lines for the He-like ions, for which all of the data are available in the literature. The HULLAC wavelengths and oscillator strengths for the $1s-np$ ($n = 2, 3$) excitations in He-like ions are compared in the table with the published values, which were calculated with perhaps more rigorous methods. For this comparison, we use primarily the compilation by Verner, Verner, & Ferland (1996) as well as values from the NIST database in cases where these are available (see Table 1). It is seen in Table 1 that the present wavelengths agree with the published values to within a few mÅ and that the oscillator strengths agree to within $\sim 10\%$ (except for Fe in Verner, Verner, & Ferland 1996). Similar comparisons for the *inner-shell* excitations in Table 1 would, of course, be even more desirable. Unfortunately, calculations for these transitions have been previously published only for Li-like Fe^{23+} . Nahar, Pradhan, & Zhang (2001) obtain for the two strongest lines of Fe^{23+} 1.860 and 1.864 Å, to be compared with the present values of 1.861 and 1.864 Å. The corresponding oscillator strengths in Nahar, Pradhan, & Zhang (2001) are 0.491 and 0.146, to be compared with the present values of 0.469 and 0.147. In order to check whether this excellent agreement is sustained for less ionized species, we have specifically calculated the $1s-2p$ lines for the Li-like ion O^{5+} , which is beyond the scope of this work and therefore not included in Table 1. Our calculations show that the two strongest $1s-2p$ lines of O^{5+} overlap at 22.01 Å and have oscillator strengths of 0.351 and 0.174. These values compare favorably with Nahar, Pradhan, & Zhang (2001), who find the lines (overlapping) at 22.05 Å, and the oscillator strengths to be 0.384 and 0.192. The agreement between the autoionization rates obtained with the two methods is of the same order as that for the oscillator strengths. In other words, even for the O^{5+} case, which could be considered an extreme upper limit to the uncertainty expected from the present values for the higher Z elements, the discrepancies with the results obtained with the more rigorous R -matrix code are not more than 0.2% for the wavelengths and 10% for the oscillator strengths and autoionization rates. The accuracy of most of the present data is, however, expected to be much higher, as indicated by the comparisons for the He-like ions and for Fe^{23+} . Having gone through this extensive comparison, it is important to stress that for an ultimate assessment of the accuracy of the atomic data, a comparison between calculations may not be sufficient. Actual measurements need to be carried out.

With the exception of maybe Fe, the lines from each ion are adequately separated in wavelength space and are readily resolvable with the gratings on board *Chandra* and *XMM-Newton*. The uncertainties associated with the wavelengths calculated with HULLAC are considerably smaller than the typical differences in wavelength. This enables an unambiguous identification of lines and, consequently, a precise measurement of the ionization state of the absorbing plasma. We note that although the inner-shell Fe lines are within the wavelength band of the HETGS spectrometer on board *Chandra*, they are difficult to resolve with HETGS. These lines will be much better detected and resolved with the microcalorimeter spectrometer (Richard et al. 1999), which is planned to fly on board *ASTRO-E2*.

For most of the inner-shell excited levels considered in this work, the total autoionization rate is much higher than the radiative decay rate, especially for the low- Z elements. The fact that the inner-shell photoexcitation processes are by and large followed by autoionization rather than X-ray re-emission is manifested in the generally low fluorescence yield, which gradually increases with the ion charge. For example, for Si and Fe, respectively, the f -weighted averages of the fluorescence yield rise gradually from ~ 0.04 and 0.75 for the F-like ion to ~ 0.27 and 0.93 for the Li-like ion. The low fluorescence yield can be associated with several important effects. First, autoionization following photoexcitation can alter the ionization balance toward higher charge states. This effect is included in ION and can be important (see § 3). Another direct result of the low fluorescence yield is that dielectronic recombination via these channels is inefficient and does not compensate for the photoexcitation-autoionization effect on the ionization balance. It further implies that these transitions are much harder to detect in emission than they are in absorption. Another consequence of the high autoionization rates is the broadening of the natural widths of the inner-shell absorption lines. At high column densities, the natural width of the line becomes important for determining the total flux absorbed.

3. SPECTRAL MODELING

We have carried out detailed spectral modeling in order to illustrate the expected intensity of the inner-shell transitions and their use as diagnostics for photoionized plasma. The calculations are performed with ION2001, the 2001 version of the photoionization code ION lately described in Netzer (1996, 2001) and Kaspi et al. (2001). The ION code computes a self-consistent model of the ionization and thermal structure of a steady state ionized gas exposed to an external radiation source. Atomic data include the most recent cross sections for the more abundant elements and all of their strong lines. Also included are f -values calculated with HULLAC for a large number of iron L-shell lines (see Kaspi et al. 2001) and dozens of M-shell iron transitions (Behar, Sako, & Kahn 2001). The new inner-shell transitions calculated in the present work are fully incorporated into the model, including their effect on the individual level population and on the level of ionization of the various elements due to photoexcitation-autoionization.

We have calculated a series of AGN-type models with a column density of 10^{22} cm^{-2} , gas density of 10^8 cm^{-3} , and a range of ionization parameters similar to the one observed in many Seyfert galaxies. The local ionizing field is defined by the spectral energy distribution (SED) of the ionizing

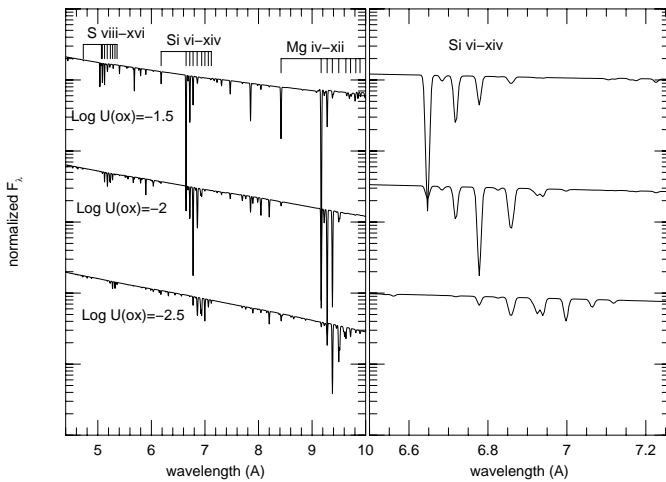


FIG. 1.—Magnesium, silicon, and sulfur $1s$ - np absorption lines for a low-density photoionized gas illuminated by a power-law X-ray continuum with $\Gamma = 1.8$ for three different ionization parameters (U_{Ox}) as indicated. A column density of 10^{22} cm^{-2} is assumed. The rulers above the spectra mark the strongest absorption lines for each of the H-like through F-like ions. *Left*: Full 4.7–10 Å range. *Right*: The silicon complex enlarged.

sources, taken here to be a single power law of photon slope $\Gamma = 1.8$ between 0.1 and 50 keV, and by the oxygen ionization parameter, U_{Ox} , defined over the 0.538–10 keV range (Netzer 2001). Guided by recent high-resolution *Chandra* observations, we have assumed a turbulent velocity of 300 km s^{-1} for the absorbing gas. Solar elemental composition is assumed in all cases.

Figure 1 shows the 4.7–10 Å wavelength range for three pure-absorption theoretical spectra calculated for U_{Ox} : $10^{-1.5}$, 10^{-2} , and $10^{-2.5}$. The largest optical-depth transitions for each of the H-like to F-like ions are marked above the plots. As evident from this example, a relatively small change in incidence flux makes a large change in the level of ionization and, consequently, in the spectral line structure. The $1s$ - $2p$ inner-shell lines of at least 2–4 L-shell ions arise for a single ionization parameter, and therefore are likely to be detected in many AGN absorption spectra. This, in combination with the observed H-like and He-like lines, provides a powerful diagnostic tool for determining very precisely the level of ionization and the elemental abundances in the absorbing gas. The strongest lines in all cases have typical equivalent widths of several mÅ and are thus easily detected in typical AGN spectra. The ample separation between lines from different charge states and the very few blends allow the identification and reliable curve-of-growth analysis for all of these lines. Thus, in addition to the ionization state, the present atomic data enable accurate column density and velocity-broadening measurements with these intrinsically narrow features. In fact, the major remaining limitation on velocity measurements is the resolution of contemporary airborne X-ray spectrometers (typically, a few 100 km s^{-1}).

The effect of the autoionization transitions following photoexcitation on the level of ionization has been tested for the $U_{\text{Ox}} = 10^{-2}$ case by running ION models with and without the inner-shell lines. We find a noticeable change in the level of ionization near the illuminated face of the cloud, where most of the absorption is expected. For example, including the inner-shell transitions changes the fractional abundance of Si^{10+} from 0.21 to 0.17, that of Si^{11+} from 0.27 to 0.25,

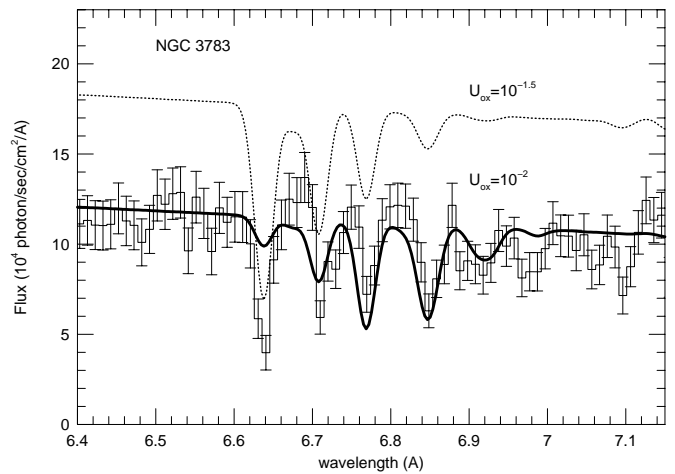


FIG. 2.—Computed inner-shell silicon lines with a turbulent velocity of 300 km s^{-1} compared with the HETGS observation of NGC 3783 (Kaspi et al. 2001). The $U_{\text{Ox}} = 10^{-2}$ model is normalized to the observed continuum and the $U_{\text{Ox}} = 10^{-1.5}$ model is arbitrarily shifted for clarity of presentation. Each absorption trough in the spectrum corresponds to one charge state of Si (compare Table 1). The feature at $\sim 7.1 \text{ Å}$ is a Mg line.

and that of Si^{12+} from 0.35 to 0.43, i.e., silicon becomes appreciably more ionized when photoexcitation-autoionization is included in the model. However, the effect on the mean level of ionization deeper in the cloud would be smaller, since the strong $1s$ - $2p$ lines saturate, which results in photoexcitation being less efficient.

The present calculations can be compared with recently obtained spectra of active galaxies. In particular, the Kaspi et al. (2001) observation of NGC 3783 is a good example of a photoionized line-of-sight gas with a line-rich absorption spectrum and with properties that are similar to the ones investigated here. In Kaspi et al. (2001), we tentatively identified features near 6.8 Å to be due to inner-shell absorption in Si. However, the atomic data available at that time did not allow for a conclusive identification. Using the newly calculated atomic data, together with our spectral code, we now revisit the Kaspi et al. (2001) observation. In particular, the present model is compared with the data in the wavelength region suspected for showing inner-shell Si lines. The results are shown in Figure 2 for two representing ionization parameters. Each absorption trough in the spectrum is attributed for the most part to a single charge state. Absorption lines of He-like Si^{12+} to N-like Si^{7+} are clearly identified in the spectrum of NGC 3783. It can be seen that the $U_{\text{Ox}} = 10^{-2}$ model reproduces the Si^{8+} - Si^{11+} lines fairly well; both the wavelengths and the equivalent widths. Figure 2 also confirms the Kaspi et al. (2001) finding that the data require a range of ionization parameters. In particular, a component with $U_{\text{Ox}} > 10^{-2}$ is required to explain the Si^{12+} line at 6.65 Å , and one with $U_{\text{Ox}} < 10^{-2}$ might be needed to account for the Si^{7+} feature at 7.00 Å . We are currently working on a more detailed fit, including lines from all of the other cosmically abundant elements to fit the new 900 ks *Chandra* HETGS data set obtained recently (Kaspi et al. 2002; George et al. 2002).

4. CONCLUSIONS

A complete set of wavelengths and oscillator strengths for inner-shell absorption lines are calculated for the entire

L-shell of Ne, Mg, Al, Si, S, Ar, Ca, and Fe. This is a continuation of our ongoing study of inner-shell absorption features in AGN spectra. The usefulness of these atomic data for plasma diagnostics is demonstrated by comparison of spectral models with the HETGS line-rich spectrum of NGC 3783. Good agreement is found between the model and the data, although for a full account of the X-ray absorber in NGC 3783, the HETGS data clearly require a multi-ionization parameter model, which will be published separately.

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REFERENCES

- Bar-Shalom, A., Klapisch, M., & Oreg, J. 2001, *J. Quant. Spectr. Radiat. Transfer*, 71, 169
- Behar, E., Sako, M., & Kahn, S. M. 2001, *ApJ*, 563, 497
- Branduardi-Raymont, G., Sako, M., Kahn, S. M., Brinkman, A. C., Kaastra, J. S., & Page, M. J. 2001, *A&A*, 365, 162
- George, I. M., et al. 2002, *ApJ*, submitted
- Kaastra, J. S., Mewe, R., Liedahl, D. A., Komossa, S., & Brinkman, A. C. 2000, *A&A*, 354, 83
- Kaspi, S., Brandt, W. N., Netzer, H., Sambruna, R., Chartas, G., Garmire, G. P., & Nousek, J. 2000, *ApJ*, 535, L17
- Kaspi, S., et al. 2001, *ApJ*, 554, 216
- . 2002, *ApJ*, submitted
- Klapisch, M., Schwob, J. L., Fraenkel, B., & Oreg, J. 1977, *J. Opt. Soc. Am.*, 67, 148
- Lee, J. L., Ogle, P. M., Canizares, C. R., Marshall, H. L., Schulz, N. S., Morales, R., Fabian, A. C., & Iwasawa, K. 2001, *ApJ*, 554, L13
- Nahar, S. N., Pradhan, A. K., & Zhang, H. L. 2001, *Phys. Rev. A*, 63, 060701(R)
- Netzer, H. 1996, *ApJ*, 473, 781
- . 2001, in *Advanced Lectures in the Starburst-AGN Connection*, ed. I. Aretxaga, D. Kunth, & R. Mujica (Singapore: World Scientific), 117
- Pradhan, A. K. 2000, *ApJ*, 545, L165
- Richard, K. L., et al. 1999, *SPIE*, 3765, 114
- Sako, M. et al. 2001, *A&A*, 365, L168
- . 2002, *ApJ*, submitted (astro-ph/0112436)
- Verner, D. A., Verner, E. M., & Ferland, G. J. 1996, *Atomic Data Nucl. Data Tables*, 64, 1